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Hot, Rotating Disks In General Relativity: Collisionless Equilibrium Models

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ABSTRACT

We present a method for constructing equilibrium disks with net angular momentum in general relativity. The method solves the relativistic Vlasov equation coupled to Einstein's equations for the gravitational field. We apply the method to construct disks that are relativistic versions of Newtonian Kalnajs disks. In Newtonian gravity these disks are analytic, and are stable against ring formation for certain ranges of their velocity dispersion. We investigate the existence of fully general relativistic equilibrium sequences for differing values of the velocity dispersion. These models are the first rotating, relativistic disk solutions of the collisionless Boltzman equation.

1. INTRODUCTION

Rotating stellar disks in dynamical equilibrium have a long history in astrophysics. Such systems are described by self-consistent solutions to the Vlasov equation for the phase-space distribution function f coupled to the equations for the gravitational field. Even in Newtonian gravitation, finding solutions is difficult since the configurations are nonspherical and have a relatively large number of nontrivial phase space degrees of freedom. (For a general review and discussion, see Fridman & Polyachenko 1984 or Binney & Tremaine 1987.) When the configuration is relativistic, the gravitational field is described by Einstein's equations and hence the problem is even more difficult.

In a previous paper (Shapiro & Teukolsky 1993; hereafter Paper I), we developed a method for constructing equilibrium axisymmetric star clusters with net rotation in general relativity. In this paper we adapt this method to treat disks. Once again we restrict our attention to the simplest phase-space distribution functions that can give rise to nonspherical equilibria, functions of particle energy E and angular momentum J_z alone. Because E and J_z are integrals of the motion, choosing a distribution function of the form $f = f(E, J_z)$ guarantees that we will have a

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solution to the Vlasov equation, provided the metric is determined self-consistently. No further dynamical equations need to be solved for the matter. By contrast, in equilibrium fluid systems one must integrate the equation of hydrostatic equilibrium.

Several researchers have studied the properties of self-gravitating *fluid* disks in general relativity. Morgan & Morgan (1969) presented an analytic model of a fluid disk with no net angular momentum. This work was followed by Bardeen & Wagoner (1969) who studied rapidly rotating, semi-analytic, zero pressure models and by Salpeter & Wagoner (1971) who looked at rapidly rotating fluid disks with some thickness. Recently the model presented by Bardeen & Wagoner (1969) was solved analytically by Meinel (1997). Collisionless disks with no net angular momentum have been modeled and evolved using a particle simulation code by Abrahams *et.al* (1994). The work presented here presents the first fully general relativistic, rotating, collisionless disk models.

As an illustration of our method, we will focus on the relativistic generalization of an important class of Newtonian disks, the Kalnajs disks (Kalnajs 1972; Binney & Tremaine 1987). These disks are completely described by simple analytic expressions. In the limiting case that all the particles move in circular orbits, the angular velocity and surface density are just those of the corresponding fluid Maclaurin spheroid in the disk limit (eccentricity $\rightarrow 1$). In Newtonian theory, however, it is known that equilibrium disks supported against collapse by rotation alone are unstable to ring formation (see, *e.g.*, Binney & Tremaine 1987, §5.3). Kalnajs (1972) showed that the disk can be stabilized by “heating” it, that is, converting some of the ordered rotational energy into random thermal motion while keeping the surface density the same. (Properties of Newtonian Kalnajs disks are reviewed in Appendix A.) We construct relativistic generalizations of Kalnajs disks and study their properties.

In addition to its potential astrophysical significance for generating models of highly relativistic stellar disks or collisionless particle distributions, the method presented here provides a class of rotating equilibria that has not been treated previously in the general relativistic literature. These collisionless models join Kerr black holes and rotating fluid stars and fluid disks as physically realistic rotating equilibria in general relativity. Such solutions to Einstein’s equations provide important insight into the effects of rotation in a strong gravitational field. They also provide useful initial data for evolution codes in general relativity. In particular, such codes can assess the stability of relativistic Kalnajs disks, since, as we will see below, we can only make heuristic statements about their stability from the results of this paper.

2. Basic Equations

In this paper we consider rotating equilibrium stellar disks that are axisymmetric. The metric can then be written in the form

$$ds^2 = -e^{\gamma+\rho} dt^2 + e^{2\sigma} (dr^2 + r^2 d\theta^2) + e^{\gamma-\rho} r^2 \sin^2 \theta (d\phi - \omega dt)^2 \quad (1)$$

where the metric potentials ρ , γ , ω , and σ are functions of r and θ only. Here and throughout we set $G = c = 1$. This is the same form of the metric used in Paper I.

All calculations in this paper are performed in the ZAMO frame (Bardeen 1970; Lightman *et al.* 1975). This is the orthonormal frame of a zero angular momentum observer, whose basis 1-forms $\omega^{\hat{\alpha}}$ are related to the coordinate basis dx^α by

$$\omega^{\hat{\alpha}} = L^{\hat{\alpha}}{}_{\alpha} dx^\alpha \quad (2)$$

where

$$L^{\hat{\alpha}}{}_{\alpha} = \begin{bmatrix} e^\nu & 0 & 0 & 0 \\ 0 & e^\nu & 0 & 0 \\ 0 & 0 & r e^\nu & 0 \\ -\omega r \sin \theta e^\beta & 0 & 0 & r \sin \theta e^\beta \end{bmatrix} \quad (3)$$

and where

$$\nu = \frac{\gamma + \rho}{2}, \quad (4)$$

$$\beta = \frac{\gamma - \rho}{2}. \quad (5)$$

The corresponding orthonormal basis vectors are given by

$$\vec{e}_{\hat{\alpha}} = L^{\alpha}{}_{\hat{\alpha}} \frac{\partial}{\partial x^\alpha} \quad (6)$$

where $L^{\alpha}{}_{\hat{\alpha}}$ is the inverse matrix of $L^{\hat{\alpha}}{}_{\alpha}$.

Following Komatsu *et al.* (1989), Cook *et al.* (1992), and Paper I, we write the Einstein field equations that determine ρ , γ and ω in the form

$$\nabla^2[\rho e^{\gamma/2}] = S_\rho(r, \mu), \quad (7)$$

$$\left(\nabla^2 + \frac{1}{r} \partial_r - \frac{\mu}{r^2} \partial_\mu \right) [\gamma e^{\gamma/2}] = S_\gamma(r, \mu), \quad (8)$$

$$\left(\nabla^2 + \frac{2}{r} \partial_r - \frac{2\mu}{r^2} \partial_\mu \right) [\omega e^{(\gamma-2\rho)/2}] = S_\omega(r, \mu), \quad (9)$$

where ∇^2 is the flat-space, spherical coordinate scalar Laplacian, $\mu = \cos \theta$, and S_ρ , S_γ and S_ω are effective source terms that include the nonlinear and matter terms. In the equations below, we explicitly exhibit the disk nature of the matter by defining a disk stress-stress energy tensor, $t_{\hat{\nu}}^{\hat{\mu}}(r)$ such that the full stress-energy tensor, $T_{\hat{\nu}}^{\hat{\mu}}(r, \mu)$ is given by

$$T_{\hat{\nu}}^{\hat{\mu}}(r, \mu) = t_{\hat{\nu}}^{\hat{\mu}}(r) \frac{\delta(\mu)}{r}. \quad (10)$$

Thus the integrals used to calculate the stress energy components, equations (31)-(33) below, need only be evaluated in the disk plane ($\mu = 0$).

The effective source terms are given by

$$S_\rho(r, \mu) \equiv R_\rho(r, \mu) + t_\rho(r) \frac{\delta(\mu)}{r} \quad (11)$$

$$S_\gamma(r, \mu) \equiv R_\gamma(r, \mu) + t_\gamma(r) \frac{\delta(\mu)}{r} \quad (12)$$

$$S_\omega(r, \mu) \equiv R_\omega(r, \mu) + t_\omega(r) \frac{\delta(\mu)}{r} \quad (13)$$

where we have defined matter dependent and matter independent source terms such that

$$R_\rho(r, \mu) = e^{\gamma/2} \left\{ \frac{1}{r} \gamma_{,r} - \frac{\mu}{r^2} \gamma_{,\mu} + \frac{\rho}{2} \left[-\gamma_{,r} \left(\frac{1}{2} \gamma_{,r} + \frac{1}{r} \right) - \frac{1}{r^2} \gamma_{,\mu} \left(\frac{1-\mu^2}{2} \gamma_{,\mu} - \mu \right) \right] \right. \\ \left. + r^2 (1-\mu^2) e^{-2\rho} \left(\omega_{,r}^2 + \frac{1-\mu^2}{r^2} \omega_{,\mu}^2 \right) \right\} \quad (14)$$

$$t_\rho(r, \mu) = 8\pi e^{\gamma/2} e^{2\sigma} \left[t_{\hat{\phi}}^{\hat{\phi}} - t_{\hat{t}}^{\hat{t}} + \frac{\rho}{2} (t_{\hat{r}}^{\hat{r}} + t_{\hat{\theta}}^{\hat{\theta}}) \right], \quad (15)$$

$$R_\gamma(r, \mu) = e^{\gamma/2} \frac{\gamma}{2} \left[-\frac{1}{2} \gamma_{,r}^2 - \frac{1-\mu^2}{2r^2} \gamma_{,\mu}^2 \right], \quad (16)$$

$$t_\gamma(r, \mu) = 8\pi e^{\gamma/2} e^{2\sigma} \left(1 + \frac{\gamma}{2} \right) (t_{\hat{r}}^{\hat{r}} + t_{\hat{\theta}}^{\hat{\theta}}), \quad (17)$$

$$R_\omega(r, \mu) = e^{(\gamma-2\rho)/2} \omega \left[-\frac{1}{r} \left(2\rho_{,r} + \frac{1}{2} \gamma_{,r} \right) + \frac{\mu}{r^2} \left(2\rho_{,\mu} + \frac{1}{2} \gamma_{,\mu} \right) + \frac{1}{4} (4\rho_{,r}^2 - \gamma_{,r}^2) \right. \\ \left. + \frac{1-\mu^2}{4r^2} (4\rho_{,\mu}^2 - \gamma_{,\mu}^2) - r^2 (1-\mu^2) e^{-2\rho} \left(\omega_{,r}^2 + \frac{1-\mu^2}{r^2} \omega_{,\mu}^2 \right) \right], \quad (18)$$

$$t_\omega(r, \mu) = 8\pi e^{(\gamma-2\rho)/2} e^{2\sigma} \left[-\omega (t_{\hat{\phi}}^{\hat{\phi}} - t_{\hat{t}}^{\hat{t}}) + \frac{\omega}{2} (t_{\hat{r}}^{\hat{r}} + t_{\hat{\theta}}^{\hat{\theta}}) - \frac{2e^\rho t_{\hat{\phi}}^{\hat{t}}}{r(1-\mu^2)^{1/2}} \right]. \quad (19)$$

Here $t_{\hat{\nu}}^{\hat{\mu}}$ are the orthonormal components of our disk stress-energy tensor for collisionless matter in the ZAMO frame (see below).

The fourth field equation determines σ and is given by

$$\sigma_{,\mu} = -\frac{1}{2} (\rho_{,\mu} + \gamma_{,\mu}) - \{ (1-\mu^2) (1+r\gamma_{,r})^2 + [\mu - (1-\mu^2)\gamma_{,\mu}]^2 \}^{-1} \\ \times \left\{ \frac{1}{2} [r^2 (\gamma_{,rr} + \gamma_{,r}^2) - (1-\mu^2) (\gamma_{,\mu\mu} + \gamma_{,\mu}^2)] [-\mu + (1-\mu^2)\gamma_{,\mu}] \right. \\ \left. + r\gamma_{,r} \left[\frac{1}{2} \mu + \mu r\gamma_{,r} + \frac{1}{2} (1-\mu^2)\gamma_{,\mu} \right] + \frac{3}{2} \gamma_{,\mu} [-\mu^2 + \mu(1-\mu^2)\gamma_{,\mu}] \right. \\ \left. - r(1-\mu^2) (\gamma_{,r\mu} + \gamma_{,r}\gamma_{,\mu}) (1+r\gamma_{,r}) - \frac{1}{4} \mu r^2 (\rho_{,r} + \gamma_{,r})^2 \right. \\ \left. - \frac{r}{2} (1-\mu^2) (\rho_{,r} + \gamma_{,r}) (\rho_{,\mu} + \gamma_{,\mu}) + \frac{1}{4} \mu (1-\mu^2) (\rho_{,\mu} + \gamma_{,\mu})^2 \right\}$$

$$\begin{aligned}
& -\frac{r^2}{2}(1-\mu^2)\gamma_{,r}(\rho_{,r}+\gamma_{,r})(\rho_{,\mu}+\gamma_{,\mu}) \\
& +\frac{1}{4}(1-\mu^2)\gamma_{,\mu}[r^2(\rho_{,r}+\gamma_{,r})^2-(1-\mu^2)(\rho_{,\mu}+\gamma_{,\mu})^2] \\
& +(1-\mu^2)e^{-2\rho}\left[\frac{1}{4}r^4\mu\omega_{,r}^2+\frac{1}{2}r^3(1-\mu^2)\omega_{,r}\omega_{,\mu}-\frac{1}{4}r^2\mu(1-\mu^2)\omega_{,\mu}^2\right. \\
& \left.+\frac{1}{2}r^4(1-\mu^2)\gamma_{,r}\omega_{,r}\omega_{,\mu}-\frac{1}{4}r^2(1-\mu^2)\gamma_{,\mu}[r^2\omega_{,r}^2-(1-\mu^2)\omega_{,\mu}^2]\right] \\
& -r^2[\mu-(1-\mu^2)\gamma_{,\mu}]e^{2\sigma}4\pi(t_{\hat{r}}-t_{\hat{\theta}})\frac{\delta(\mu)}{r} \\
& +r^2(1-\mu^2)^{1/2}(1+r\gamma_{,r})e^{2\sigma}8\pi t_{\hat{\theta}}\frac{\delta(\mu)}{r}\Big\}. \tag{20}
\end{aligned}$$

The terms containing $t_{\hat{\theta}}$ and $t_{\hat{r}}$ in the above equation are identically zero for an equatorial disk since each particle has no θ -component of momentum. The remaining matter term in equation (20),

$$-r^2[\mu-(1-\mu^2)\gamma_{,\mu}]e^{2\sigma}4\pi t_{\hat{r}}\frac{\delta(\mu)}{r} \tag{21}$$

also does not contribute since $\gamma_{,\mu}|_{\mu=0}$ vanishes by symmetry.

The stress-energy tensor for the matter is determined by the phase-space distribution function f , which is governed by the relativistic Vlasov equation (the matter dynamical equation). Any distribution function of the form $f = f(E, J_z)$ is an equilibrium solution of the Vlasov equation in axisymmetry. Here E and J_z are the two constants of motion associated with the Killing vectors $\partial/\partial t$ and $\partial/\partial\phi$. For our two-dimensional disk system they are

$$E \equiv -\mathbf{p} \cdot \frac{\partial}{\partial t} = e^{\nu_e} p^{\hat{t}} + \omega e^{\beta_e} r p^{\hat{\phi}}, \tag{22}$$

$$J_z \equiv \mathbf{p} \cdot \frac{\partial}{\partial \phi} = e^{\beta_e} r p^{\hat{\phi}}, \tag{23}$$

where $p^{\hat{\alpha}}$ are the orthonormal components of the particle 4-momentum \mathbf{p} in the ZAMO frame and we have defined, for convenience, the quantities,

$$\begin{aligned}
\nu_e & \equiv \nu(r, 0), \\
\beta_e & \equiv \beta(r, 0), \\
\omega_e & \equiv \omega(r, 0), \quad etc.
\end{aligned} \tag{24}$$

Physically, E is the conserved energy of a particle and J_z is the conserved angular momentum about the symmetry axis.

Working in two-dimensions, we define the disk stress-energy tensor for the matter,

$$t^{\hat{\alpha}\hat{\beta}} = \int f p^{\hat{\alpha}} p^{\hat{\beta}} \frac{d^2 \hat{p}}{p^{\hat{t}}}. \tag{25}$$

Here, for a disk, we have

$$p^{\hat{t}} = [(p^{\hat{r}})^2 + (p^{\hat{\phi}})^2 + m^2]^{1/2}, \quad (26)$$

and

$$d^2\hat{p} = dp^{\hat{r}} dp^{\hat{\phi}}, \quad (27)$$

where m is the particle mass. For the field equations (7) – (9) and (20), we need only the combinations

$$t_1 \equiv t_{\hat{\phi}}^{\hat{\phi}} - t_{\hat{t}}^{\hat{t}}, \quad (28)$$

$$t_2 \equiv t_{\hat{r}}^{\hat{r}}, \quad (29)$$

$$t_{03} \equiv t_{\hat{\phi}}^{\hat{t}}. \quad (30)$$

Using equations (25) – (27), we get

$$t_1(r) = \int dp^{\hat{r}} \int dp^{\hat{\phi}} \frac{(p^{\hat{r}})^2 + 2(p^{\hat{\phi}})^2 + m^2}{[(p^{\hat{r}})^2 + (p^{\hat{\phi}})^2 + m^2]^{1/2}} f(E, J_z), \quad (31)$$

$$t_2(r) = \int dp^{\hat{r}} \int dp^{\hat{\phi}} \frac{(p^{\hat{r}})^2}{[(p^{\hat{r}})^2 + (p^{\hat{\phi}})^2 + m^2]^{1/2}} f(E, J_z), \quad (32)$$

$$t_{03}(r) = \int dp^{\hat{r}} \int dp^{\hat{\phi}} p^{\hat{\phi}} f(E, J_z). \quad (33)$$

At each spatial point (r, μ) , the limits of integration in equations (31) – (33) will be determined by the distribution function, as discussed in the next subsection.

2.1. Distribution Function

The Newtonian Kalnajs distribution function is given by equation (A4). We construct a relativistic generalization by letting $E_N \rightarrow E - m$:

$$f_{GR}(E, J_z) = \begin{cases} K \left[2\{(bV J_z/R_m) - (E - m)\} - V^2(1 + b^2) \right]^{-1/2} & [\dots] > 0 \\ 0 & [\dots] \leq 0 \end{cases} \quad (34)$$

Here V , b and R_m are constants which in the Newtonian limit have the following interpretations:

$$b \equiv \frac{\text{Mean angular rotation rate}}{\text{Angular speed of a circular orbit}} \equiv \frac{\Omega}{\Omega_{circ}} \equiv \frac{\Omega R_m}{V}, \quad (35)$$

and R_m is the matter radius. Note that in the Newtonian limit we have

$$E - m \rightarrow E_N \quad J_z \rightarrow J_{zN}, \quad (36)$$

where E_N and J_{zN} are the Newtonian energy and angular momentum as defined by equations (A5) and (A6). Thus we recover the Newtonian distribution function f_N , equation (A4), from equation (34) in the Newtonian limit.

Since the distribution function is nonzero only for positive values of the argument of the square root, we can solve for the limits of integration for the stress-energy integrals (eqs. (31)–(33)) by using the expressions (22) and (23) for E and J_z in equation (34) and setting the argument of the square root to zero. Thus we find the limits to be

$$0 < p^{\hat{r}} < p_{max}^{\hat{r}}(p^{\hat{\phi}}, r) \quad (37)$$

and

$$p_{-}^{\hat{\phi}}(r) < p^{\hat{\phi}} < p_{+}^{\hat{\phi}}(r) \quad (38)$$

with

$$p_{max}^{\hat{r}}(p^{\hat{\phi}}, r) \equiv \left[\frac{a^2 + d^2 - 1}{1 - d^2} - (1 - d^2) \left(p^{\hat{\phi}} - \frac{ad}{1 - d^2} \right)^2 \right]^{1/2} \quad (39)$$

and

$$p_{\pm}^{\hat{\phi}} \equiv \frac{ad}{1 - d^2} \pm \frac{(a^2 + d^2 - 1)^{1/2}}{1 - d^2}. \quad (40)$$

Here a and d are functions of r only and are given by

$$a(r) \equiv e^{-\nu_e(r)} \left[1 - \frac{V^2}{2}(1 - b^2) \right] \quad (41)$$

and

$$d(r) \equiv e^{\beta_e(r) - \nu_e(r)} r \left[b \frac{V}{R_m} - \omega_e(r) \right]. \quad (42)$$

Note that in Paper I the distribution functions considered contained a parameter E_{max} that was less than m for bound systems of finite extent. The limits of integration were determined by $E \leq E_{max}$. Here, one can verify that the limits (37)–(38) never violate the condition $E \leq m$. Curiously, equality can be attained even in the Newtonian limit: $E_N \rightarrow 0$ at the outer edge of a “cold” Kalnajs disk.

The total mass-energy of the system is

$$\begin{aligned} M &= - \int (2T_{\nu}^{\mu} - \delta_{\nu}^{\mu} T) \xi_{(t)}^{\nu} d^3 \Sigma_{\mu} \\ &= \int (-2T_t^t + T) \sqrt{-g} d^3 x. \end{aligned} \quad (43)$$

Here, $\xi_{(t)}^{\nu} = \partial/\partial t$ is the time Killing vector. Transforming the integrand to the ZAMO frame, and integrating out the δ -function in μ , we obtain

$$M = 2\pi \int [t_1(r) + t_2(r) + 2\omega_e e^{-\rho_e} r t_{03}(r)] e^{2\sigma_e + \gamma_e} r dr. \quad (44)$$

As we will see in Section 5, this equation plays a crucial role in the iterative solution of the combined matter and field equations.

3. Diagnostic Probes

There are a number of useful quantities that characterize an equilibrium system once a solution has been obtained. The total angular momentum is given by

$$\begin{aligned} J &= \int T_{\nu}^{\mu} \xi_{(\phi)}^{\nu} d^3 \Sigma_{\mu} \\ &= \int T_{\phi}^t \sqrt{-g} d^3 x. \end{aligned} \quad (45)$$

Transforming to the ZAMO frame, and again integrating out the delta function in μ , we get

$$J = 2\pi \int t_{03}(r) e^{2(\sigma_e + \beta_e)} r^2 dr. \quad (46)$$

The surface rest mass density is given by

$$\sigma_0(r) = m \int dp^{\hat{r}} \int dp^{\hat{\phi}} f_{GR}(E, J_z), \quad (47)$$

and the total rest mass by

$$M_0 = 2\pi \int \sigma_0(r) e^{2\sigma_e + \beta_e} r dr \quad (48)$$

The binding energy of the system is defined as

$$E_b \equiv M_0 - M, \quad (49)$$

and can be computed from equations (44) and (48).

4. Scaling and Non-dimensional Units

The quantities m and M can be scaled out of all the above equations. For example, we can define

$$\begin{aligned} \tilde{p} &= p/m, \quad \tilde{r} = r/M, \quad \tilde{J} = J/M^2 \\ \tilde{t}^{\alpha\beta} &= t^{\alpha\beta}/M^{-1}, \quad \tilde{f}_{GR} = f_{GR}/M^{-1}m^{-3}, \quad \tilde{J}_z = J_z/mM. \end{aligned} \quad (50)$$

With these definitions, all the previous equations can be written in tilde variables without m or M appearing. Equivalently, the original equations can be solved setting $m = M = 1$ and scaling the final results according to equation (50) to accommodate arbitrary values of m and M . Henceforth we will make this simplification.

In the Newtonian limit there is an additional scale freedom in that we can also set R_m/M to one. This is not the case for relativistic systems.

5. Numerical Solution

The numerical scheme used here is a 2-dimensional analogue of the procedure adopted in Paper I. In this procedure we start with initial guesses for the metric potentials, ρ , γ , ω , and σ (see section 6 below). We then integrate equations (31)-(33) to find t_1 , t_2 and t_{03} up to the constant factor K appearing in equation (34). This unknown factor K is fixed by requiring that the total mass of the system, equation (44), satisfy $M = 1$. Next we integrate the field equations (7)-(8) for ρ , γ , and ω , and then equation (20) for σ . We then iterate this procedure until some convergence criterion is met.

Using a combination of integral and finite differencing techniques we solve the equations for the matter and gravitational fields. These equations are solved on a discrete grid in μ and r on the computational domain $0 \leq r \leq \infty$ and $0 \leq \mu \leq 1$. Unlike the Newtonian case, we cannot restrict the computational domain to the matter interior because the matter-independent effective source terms R_ρ , R_γ and R_ω are nonzero in the vacuum exterior. Consequently we divide the radial grid into an interior and exterior domain. Each domain is covered by a geometrically spaced grid in r , with the grids joined smoothly at the cluster surface. We use an angular grid that is uniformly spaced in μ . The interior radial grid is arranged to yield sufficient resolution for the core of the cluster, while the outer grid extends to some sufficiently large radius, typically 3 times the radius of the matter surface. High resolution of the core is crucial for obtaining numerical accuracy in highly centrally condensed relativistic disks.

The three elliptic field equations (7) – (9) are solved by an integral Green’s function approach following Komatsu *et al.* (1989) and Cook *et al.* (1992). Again we make a matter-dependent, matter-independent split of the Green’s functions depending on whether we are integrating the matter-dependent source terms t_ρ , t_γ , and t_ω or the matter-independent source terms R_ρ , R_γ and R_ω . Thus we write the solution of equation (7) as

$$\rho(r, \mu) = W_\rho(r, \mu) + G_\rho(r, \mu) \quad (51)$$

where we have defined

$$W_\rho(r, \mu) \equiv - \sum_{n=0}^{\infty} e^{-\gamma/2} \int_0^\infty dr' \int_0^1 d\mu' r'^2 f_{2n}^2(r, r') P_{2n}(\mu) P_{2n}(\mu') R_\rho(r', \mu'), \quad (52)$$

and

$$G_\rho(r, \mu) \equiv - \frac{1}{2} \sum_{n=0}^{\infty} e^{-\gamma/2} P_{2n}(\mu) \frac{(-1)^n (2n-1)!!}{2^n n!} \int_0^\infty dr' r' f_{2n}^2(r, r') t_\rho(r', 0). \quad (53)$$

Here we have used the δ -function in μ to carry out the μ' integration in G_ρ . Similarly to solve equation (8) we write

$$r \sin \theta \gamma(r, \mu) = W_\gamma(r, \mu) + G_\gamma(r, \mu) \quad (54)$$

with

$$W_\gamma(r, \mu) \equiv - \frac{2}{\pi} \sum_{n=1}^{\infty} e^{-\gamma/2} \int_0^\infty dr' \int_0^1 d\mu' r'^2 f_{2n-1}^1(r, r') \frac{1}{2n-1} \sin(2n-1)\theta$$

$$\times \sin(2n-1)\theta' R_\gamma(r', \mu'), \quad (55)$$

and

$$G_\gamma(r, \mu) \equiv \frac{1}{\pi} \sum_{n=1}^{\infty} e^{-\gamma/2} \frac{\sin(2n-1)\theta(-1)^n}{(2n-1)} \int_0^\infty dr' f_{2n-1}^1(r, r') t_\gamma(r', 0). \quad (56)$$

Finally, to solve equation (9) we write

$$r \sin \theta \omega(r, \mu) = W_\omega(r, \mu) + G_\omega(r, \mu) \quad (57)$$

with

$$\begin{aligned} W_\omega(r, \mu) \equiv & - \sum_{n=1}^{\infty} e^{(2\rho-\gamma)/2} \int_0^\infty dr' \int_0^1 d\mu' r'^3 \sin \theta' f_{2n-1}^2(r, r') \frac{1}{2n(2n-1)} \\ & \times P_{2n-1}^1(\mu) P_{2n-1}^1(\mu') R_\omega(r', \mu'), \end{aligned} \quad (58)$$

and

$$G_\omega(r, \mu) \equiv -\frac{1}{2} \sum_{n=1}^{\infty} e^{(2\rho-\gamma)/2} \frac{P_{2n-1}^1(\mu)}{2n(2n-1)} \frac{(-1)^n (2n-1)!!}{2^n (n-1)!} \int_0^\infty dr' r'^2 f_{2n-1}^2(r, r') t_\omega(r', 0). \quad (59)$$

Here

$$f_n^1(r, r') = \begin{cases} (r'/r)^n, & \text{for } r'/r \leq 1 \\ (r/r')^n, & \text{for } r'/r > 1 \end{cases} \quad (60)$$

and

$$f_n^2(r, r') = \begin{cases} (1/r)(r'/r)^n, & \text{for } r'/r \leq 1, \\ (1/r')(r/r')^n, & \text{for } r'/r > 1 \end{cases} \quad (61)$$

Among the advantages of this Green's function approach for solving the elliptic field equations is that the asymptotic conditions on ρ , γ and ω are imposed automatically. That is, $\rho \sim \mathcal{O}(1/r)$, $\gamma \sim \mathcal{O}(1/r^2)$, and $\omega \sim \mathcal{O}(1/r^3)$ for large r . To improve the accuracy of the angular integrations, we use the identities in equations (34) – (38) of Cook *et al.* (1992).

It should be noted that in the integration of equation (59), the calculation of $t_\omega(r, 0)$ at $r = 0$ requires special care. It can be seen from equations (33), (39) and (40) that $t_{03}(r = 0) = 0$, while regularity conditions near the axis imply that $t_{03} \sim r$ near $r = 0$. Thus, at $r = 0$, the calculation of the quantity t_{03}/r that appears in equation (19) is done by computing the quantity $dt_{03}/dr|_{r=0}$ analytically. We do this by taking a derivative of equation (33) directly. For any r this gives

$$dt_{03}/dr = \frac{d}{dr} \left[\int_{p_-^{\hat{\phi}}(r)}^{p_+^{\hat{\phi}}(r)} dp^{\hat{\phi}} p^{\hat{\phi}} \int_0^{p_{max}^{\hat{\phi}}(r, p^{\hat{\phi}})} dp^{\hat{r}} f_{GR}(E, J_z) \right] \quad (62)$$

$$\begin{aligned} &= \int_{p_-^{\hat{\phi}}(r)}^{p_+^{\hat{\phi}}(r)} dp^{\hat{\phi}} p^{\hat{\phi}} \frac{\partial I(p^{\hat{\phi}}, r)}{\partial r} \\ &\quad + p_+^{\hat{\phi}}(r) I(p_+^{\hat{\phi}}(r), r) \frac{dp_+^{\hat{\phi}}(r)}{dr} - p_-^{\hat{\phi}}(r) I(p_-^{\hat{\phi}}(r), r) \frac{dp_-^{\hat{\phi}}(r)}{dr}. \end{aligned} \quad (63)$$

Here $p_{\pm}^{\hat{\phi}}(r)$ and $p_{max}^{\hat{\phi}}(r, p^{\hat{\phi}})$ are defined in equations (40) and (39) and we have defined $I(p^{\hat{\phi}}, r)$ to be

$$I(p^{\hat{\phi}}, r) = \int_0^{p_{max}^{\hat{\phi}}(r, p^{\hat{\phi}})} dp^{\hat{\phi}} f_{GR}(E, J_z). \quad (64)$$

At $r = 0$ it can be shown that the first term in equation (63) vanishes by symmetry. The other two terms involve the integral $I(p^{\hat{\phi}}, r)$ evaluated at $p^{\hat{\phi}} = p_{\pm}^{\hat{\phi}}(0)$ and $r = 0$. From equation (64) and the form of the distribution function f_{GR} (equation (34)) we see that this integrand is divergent and the limits collapse to zero at $p^{\hat{\phi}} = p_{\pm}^{\hat{\phi}}(0)$. However, near $p^{\hat{\phi}} = p_{\pm}^{\hat{\phi}}(0)$ we can expand the integrand (and upper limit) to get an estimate of the value of the integral for use in the calculation of $dt_{03}/dr|_{r=0}$. This yields

$$dt_{03}/dr|_{r=0} = \pi \left[\sqrt{a(r)(a^2(r) - 1)} a(r) d_r(r) e^{-\nu_e/2} \right]_{r=0}. \quad (65)$$

Here $a(r)$ and $d(r)$ are functions defined by equations (41) and (42).

Equation (20) for σ is solved by integrating the linear ODE from the pole ($\mu = 1$) to the equator with the initial condition that

$$\sigma = \frac{\gamma - \rho}{2} \quad \text{at} \quad \mu = 1, \quad (66)$$

which arises from the requirement of local flatness on the coordinate axis. The derivatives of ρ , γ and ω appearing in the matter-independent source terms R_{ρ} , R_{γ} , R_{ω} and the right-hand side of equation (20) are evaluated by finite differencing.

Since we can evaluate the integrands in equations (31) – (33) at any values of $p^{\hat{\phi}}$ and $p^{\hat{\phi}}$, we carry out the quadrature over each variable by Romberg integration (Press *et al.* 1992).

For a convergent solution, we require the maximum fractional change in all four metric functions to be less than 1×10^{-3} on successive iterations. This typically takes about 20 iterations when we start with a Newtonian initial guess (Section 6). A significant number of iterations can be saved during a sequence of calculations in which the heating parameter b is varied if we use the previous solution as the initial guess for the next value of b .

For disks in the relativistic region, 600 radial zones, 20 angular zones, and 20 Legendre polynomials are adequate.

6. Initial Data

As an initial guess for the metric potentials $\rho(r, \mu)$, $\gamma(r, \mu)$, $\omega(r, \mu)$, $\sigma(r, \mu)$ we used their Newtonian limits, given in terms of the Newtonian potential Φ_N , by

$$\rho(r, \mu) \rightarrow 2\Phi_N(r, \mu),$$

$$\begin{aligned}\gamma(r, \mu) &\rightarrow 0, \\ \sigma(r, \mu) &\rightarrow -\Phi_N(r, \mu), \\ \omega(r, \mu) &\rightarrow 0.\end{aligned}\tag{67}$$

Mihalas (1969) gives an expression for the Newtonian potential of an oblate homogeneous spheroid. By taking the limit as the eccentricity $e \rightarrow 1$, we obtain the Newtonian potential $\Phi_N(r, \mu)$ for all points both on and off the disk. Outside the matter, we find

$$\Phi_N(r, \mu) = -\frac{3M}{2R_m} \left[B - \frac{q_s}{2(1-x^2)^{1/2}} \left((1-3\mu^2)(B(1-x^2)^{1/2} - x) + (1-\mu^2)x^3 \right) \right]. \tag{68}$$

where $q_s \equiv r^2/R_m^2$, $x = \sqrt{p}$, and $B = \arcsin x$ with p defined by the quadratic equation

$$q_s(1-\mu^2)p^2 - (q_s+1)p + 1 = 0. \tag{69}$$

Inside the matter, we have

$$\Phi_N(r, \mu) = -\frac{3\pi M}{4R_m} \left(1 - \frac{q_s}{2} \right). \tag{70}$$

7. Numerical Results

Our region of investigation spans values of the heating parameter in the range $.05 \leq b \leq .95$ for values of R_m/M ranging from $R_m/M = 455$ down to $R_m/M = 6.26$. Figure 1 (a) shows the convergence of some representative runs in the Newtonian region. As we can see, as b increases, R_m/M of the solution decreases. This is also evident for the more relativistic cases shown in Figure 1 (b).

To better understand our solutions it is helpful to look at how the central redshift, z_c , varies with the heating parameter, b . For the Newtonian Kalnajs disk given in Appendix A, we can calculate the central redshift, z_c analytically from the Newtonian gravitational potential as follows:

$$\begin{aligned}z_c &\equiv [-g_{tt}(0,0)]^{-1/2} - 1 \\ &= \left[e^{\gamma(0,0)+\rho(0,0)} \right]^{-1/2} - 1.\end{aligned}\tag{71}$$

Now use equations (67) to get

$$z_c^{Newtonian} = -\phi_N(0,0) = \frac{3\pi M}{4R}. \tag{72}$$

This quantity is independent of the heating parameter b and thus is a useful diagnostic. Although analytic equilibrium solutions exist in the Newtonian case even in the unstable region ($b \geq .816$) (see Appendix A), it is not obvious that our iterative method will converge to them. Instabilities may be signaled by failure to converge, or by causing convergence to a solution that is far from the true Kalnajs solution. Thus by plotting the central redshift versus the heating parameter we can

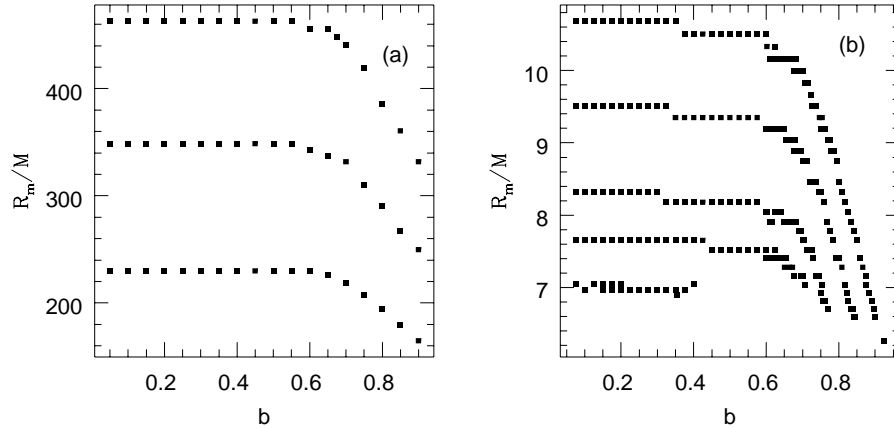


Fig. 1.— Convergent runs for some representative values of R_m/M . (a) Runs in the Newtonian region. (b) Runs in the more relativistic region.

hope to see where the solution we find differs from the true Newtonian solution. As can be seen in Figure 2 (a), for a Newtonian cluster with $R_m/M = 455$, our code finds the Kalnajs equilibrium solutions for $b \leq .6$. Since the stability region is given by $0 \leq b < .816$, the change in the central redshift does not serve as a sharp indicator of the transition to the unstable region.

Similar behavior can be observed in our relativistic clusters. Figure 2 (b) shows an example for $R_m/M = 8.84$. Regions in which z_c is not constant may represent true general relativistic equilibria, but most likely these clusters are near or inside the general relativistic unstable region. To fully pin down this region requires a dynamical stability analysis, which can probably only be done by numerical evolution.

The most relativistic solution we have found has an R_m/M of 6.97. By comparing this disk to the Kerr geometry we can get an idea of how relativistic this model is. Using the proper circumference of the disk we find that our most relativistic cluster corresponds to an R/M for Kerr of 8.09. Table 1 gives some properties of the representative models discussed here.

In Figure 3 we show contour plots of the surface density for two examples of our equilibrium disks. Frame (a) shows our most Newtonian cluster ($R_m/M = 455$) with $b = .6$. In Newtonian theory these disks have homogeneous volume density. Frame (b) shows the same plot for the analytic version of the same cluster. In contrast, Frames (c) and (d) show our most relativistic cluster ($R_m/M = 6.97$) with $b = .1$ and its Newtonian counterpart respectively. The relativistic disks are in general more centrally condensed than their Newtonian cousins.

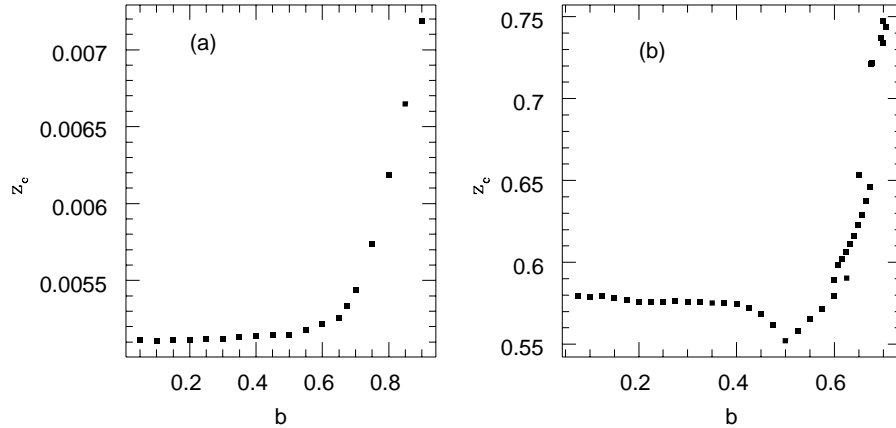


Fig. 2.— Central redshift (z_c) versus heating parameter b for our most Newtonian cluster, (a) $R_m/M = 455$ and a representative relativistic cluster, (b) $R_m/M = 8.84$.

In previous work (Shapiro & Teukolsky 1985,1992b) on polytropic axisymmetric clusters we found that a maximum in the binding energy along an equilibrium sequence marked the onset of a dynamical instability. Since, for the Newtonian Kalnajs disks here the potential, (equation (A7)) and surface matter density, (equation (A1)) are not functions of the heating parameter, the binding energy is also not a function of b . Thus it cannot be used to analyze stability. However, in the general relativistic case we expect this degeneracy to be lifted. We do indeed find that the binding energy does depend on b . However, because of numerical errors, we are not able to reliably find the turning point in the binding energy that would signal the onset of instability. Again, a firm limit on the stability region must wait for a full dynamical evolution code.

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A. Newtonian Kalnajs Disks

A Newtonian Kalnajs disk can be obtained by “flattening” a homogeneous oblate spheroid, i.e., by letting the eccentricity $e \rightarrow 1$ (Binney & Tremaine 1987). This flattening yields a disk with

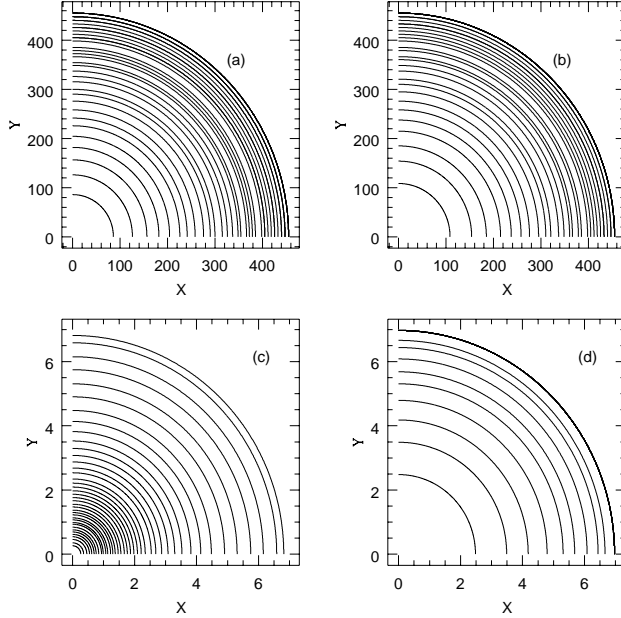


Fig. 3.— Surface density contour plots: (a) Our most Newtonian cluster with $R_m/M = 455$. and (b) its analytic analogue. (c) Our most relativistic cluster with $R_m/M = 6.97$. (d) The analytic Newtonian version of this cluster ($R_m/M = 6.97$).

a surface density given by

$$\Sigma(R) = \Sigma_c(1 - R^2/R_m^2), \quad (\text{A1})$$

where Σ_c is the central density, R_m is the radius of the disk (matter radius) and R is the radius in the disk plane defined by

$$R \equiv r \sin \theta. \quad (\text{A2})$$

When all the particles are in circular orbits with uniform angular velocity Ω_{circ} , then Ω_{circ} is related to the central density by the same relation as for Maclaurin spheroids.

Kalnajs (1972) developed a family of equilibrium disks, all of which have the above surface density. Each member of this family is characterized by the parameter b defined in equation (35).

These disks have velocity dispersions that are governed by the tunable parameter b :

$$\overline{(v^\phi)^2} - \left(\overline{v^\phi}\right)^2 = \overline{(v^{\hat{R}})^2} = \frac{1}{3}V^2(1 - b^2)(1 - R^2/R_m^2). \quad (\text{A3})$$

From the above equation we can see that when $b \approx 1$ ($\Omega R_m \approx V$) these disks are “cold”, i.e. have very little thermal motion. This corresponds physically to a system in which all particles move on nearly circular orbits. On the other hand, when $b \ll 1$ ($\Omega R_m \ll V$) we have “hot”

systems in which most of the support against self-gravity comes from the random motions given in equation (A3) (Binney & Tremaine 1987).

A distribution functions that generates this family is

$$f_N(E_N, J_{zN}) = \begin{cases} K_N \left[2\{(bV J_{zN}/R_m) - E_N\} - V^2(1 + b^2) \right]^{-1/2} & [\dots] > 0 \\ 0 & [\dots] \leq 0 \end{cases} \quad (\text{A4})$$

Here K_N is a constant that can be determined by fixing the total mass of the system, E_N is the Newtonian energy, and J_{zN} is the Newtonian angular momentum about the symmetry (z) axis. These quantities, with $m = 1$, are given by

$$E_N = \frac{(p^{\hat{R}})^2}{2} + \frac{(p^{\hat{\phi}})^2}{2} + \phi_N(R) \quad (\text{A5})$$

and

$$J_{zN} = R p^{\hat{\phi}}. \quad (\text{A6})$$

The Newtonian potential, $\phi_N(R)$, in the disk plane is

$$\phi_N(R) \equiv \Phi_N(r, 0) = -V^2 \left(1 - \frac{R^2}{2R_m^2} \right). \quad (\text{A7})$$

Using the linearized collisionless Boltzmann equation (linearized Vlasov equation) Kalnajs (1972) analyzed the stability of this family of disks. He found that restricting the heating parameter, b , to be less than 0.816 produced disk systems that are stable against all axisymmetric disturbances, i.e. disks with $0 \leq b < .816$ are stable against ring formation.

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Table 1. Properties of some representative solutions.

R_m/M	b	z_c	g_{tt}^{max} ^a	E_b ^b	J ^c
455	0.6	0.0051	-0.9897	-0.0032	4.2304
6.97	0.1	.9556	-0.2614	... ^d	0.0693

^a g_{tt}^{max} is the maximum value of g_{tt} .

^b E_b is in units of M_0 .

^c J is in units of M^2

^dFor the relativistic disk here our value of E_b is unreliable because of numerical errors.